

НАБЛЮДЕНИЯ МЕЛКОМАСШТАБНЫХ ЭНЕРГЕТИЧЕСКИХ СОБЫТИЙ В СОЛНЕЧНОЙ ПЕРЕХОДНОЙ ОБЛАСТИ: ВЗРЫВНЫХ СОБЫТИЙ, ВСПЛЕСКОВ УЛЬТРАФИОЛЕТОВОГО ИЗЛУЧЕНИЯ И СТРУЙНЫХ ЯВЛЕНИЙ

OBSERVATIONS OF SMALL-SCALE ENERGETIC EVENTS IN THE SOLAR TRANSITION REGION: EXPLOSIVE EVENTS, UV BURSTS, AND NETWORK JETS

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Аннотация. В данной статье рассматриваются проблемы наблюдения трех общих мелкомасштабных энергетических событий в солнечной переходной области, а именно, взрывные события, всплески ультрафиолетового излучения и струйные явления. Эти события определяются по спектральным данным или (и) по изображениям. Разработка многочисленных инструментов, которые могут использоваться для наблюдений переходной области, позволила в последние годы получить обширную информацию об этих явлениях. Эти события дают примерное представление о переносе массы и энергии между солнечной хромосферой и короной. Поскольку физические механизмы, ответственные за эти мелкомасштабные события, могут быть схожи с механизмами, ответственными за крупномасштабные явления, такие как вспышки и выброс корональной массы, анализ этих событий мог бы также способствовать пониманию мелкомасштабных и крупномасштабных процессов в солнечной атмосфере. Наблюдения этих мелкомасштабных энергетических событий показывают, что переходная область чрезвычайно динамична и является слоем солнечной атмосферы, определяющим процессы, протекающие между хромосферой и короной.

Ключевые слова: переходная область, магнитное пересоединение, мелкомасштабная динамика, спектроскопические наблюдения.

Abstract. In this paper, we review observational aspects of three common small-scale energetic events in the solar transition region (TR), namely TR explosive events, ultraviolet bursts and jets. These events are defined in either (both) spectral or (and) imaging data. The development of multiple instruments capable of observing the TR has allowed researchers to gain numerous insights into these phenomena in recent years. These events have provided a proxy to study how mass and energy are transported between the solar chromosphere and the corona. As the physical mechanisms responsible for these small-scale events might be similar to the mechanisms responsible for large-scale phenomena, such as flares and coronal mass ejections, analysis of these events could also help our understanding of the solar atmosphere from small to large scales. The observations of these small-scale energetic events demonstrate that the TR is extremely dynamic and is a crucial layer in the solar atmosphere between the chromosphere and the corona.

Keywords: transition region, magnetic reconnection, small-scale dynamics, spectroscopic observations.

THE TRANSITION REGION OF THE SOLAR ATMOSPHERE

The transition region (TR) of the solar atmosphere is a region between the solar chromosphere and corona, which is assumed to be a thin layer (a few hundred kilometers) in one-dimensional static models of the solar atmosphere (e.g. [Vernazza et al., 1981; Avrett, Loeser, 2008]). Across this region, the temperature of the solar atmosphere rises from about $2 \cdot 10^4$ K in the upper chromosphere to about 10^6 K in the corona [Mariska, 1992].

As a critical region between the chromosphere and corona, TR holds the key to understanding how energy and mass are transported from the lower to higher solar atmosphere. It therefore attracts a lot of attention of the community. Recent observational and theoretical studies have revealed that the TR is highly dynamic and nonuniform. Figure 1 displays schematic structures of the lower solar atmosphere, which presents dynamics and complexity in the region. Thus, the TR is often regarded as a temperature region rather than a real geometrical layer. More details (both historical and recent advance)

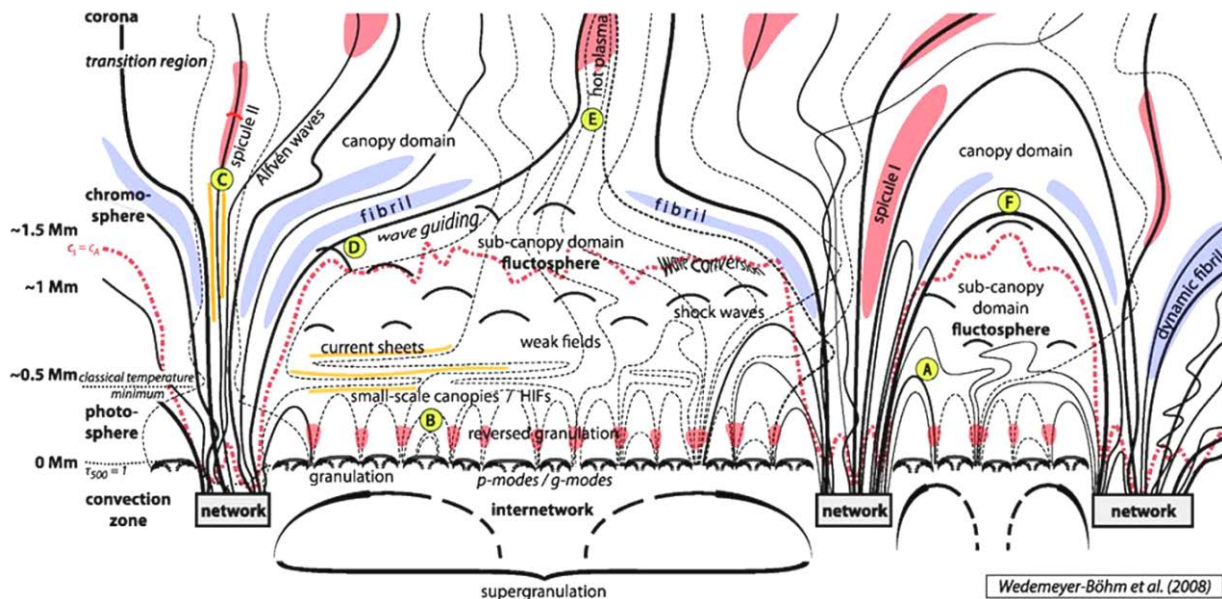


Figure 1. Schematic structures of the lower solar atmosphere (adapted from [Wedemeyer-Böhm, 2009])

of the TR could be found in review articles [Mariska, 1986, 1992; Tian, 2017; Young et al., 2018].

TR temperatures show that the TR mostly emits in far ultraviolet (FUV) especially by ions such as C IV, O IV and Si IV. Thus, observations of the solar transition region became available only in the space era. The Naval Research Laboratory’s High Resolution Telescope and Spectrograph (HRTS) [Bartoe, Brueckner, 1975] recorded ultraviolet spectra of the Sun during 10 shuttle flights, which revealed many dynamic phenomena in the spectral lines formed in the TR (see e.g. [Brueckner, Bartoe 1983; Dere et al., 1989, 1991]. Later in the 1990s, the Solar and Heliospheric Observatory (SOHO) [Domingo et al., 1995] was launched carrying many instruments including two that could observe the TR, the Solar Ultraviolet Measurements of Emitted Radiation (SUMER) [Wilhelm et al., 1995; Lemaire et al., 1997] and the Coronal Diagnostic Spectrometer (CDS) [Harrison et al., 1995]. SUMER provided full wavelength spectra of the Sun from 500 to 1610 Å, which have formation temperatures from 10^4 to 10^6 K. SUMER is a powerful tool for studies of the TR, and has revealed various small-scale structures of the TR in terms of intensity, velocity, density, and temperature (see, e.g., [Hassler et al., 1999; Xia et al., 2003, 2004; Tu et al., 2005]. CDS is designed to detect solar extreme ultraviolet radiation that allows us to examine conditions in the solar corona. The spectra recorded by CDS also include a few TR lines that help to discover some small-scale dynamics such as blinkers [Harrison, 1997] in the region. The Transition Region and Coronal Explorer (TRACE) [Handy et al., 1999], launched in late 1990s, provided us with high spatial-temporal resolution images of the solar atmosphere that allowed us to follow the evolution of the dynamic phenomena in the TR. A few more satellites aiming at studying the Sun were launched in the first 20 years of the 21st century. An important mission is Hinode [Kosugi et al., 2007] launched in 2006. Hinode carries three instruments: the Extreme Ultraviolet Imaging Spectrometer (EIS) [Culhane et al., 2007], Solar Optical Telescope (SOT) [Tsuneta et al., 2008], and X-Ray Telescope (XRT)

[Golub et al., 2007], which could take spectra, X-ray images, and magnetic field data of the Sun. These data could help understand the connection among activities seen in different temperatures from that in the photosphere to that in the corona. Another very successful mission to study the Sun is the Solar Dynamics Observatory (SDO) [Pesnell et al., 2012] launched in 2010. Aboard SDO there are two instruments: the Atmospheric Imaging Assembly (AIA) [Lemen et al., 2012] and the Helioseismic and Magnetic Imager (HMI) [Scherrer et al., 2012], which take high spatial-temporal resolution images of the full-disk. AIA takes continuous full-disk observations of the solar chromosphere and corona in seven extreme ultraviolet (EUV) channels, spanning a temperature range from approximately $20 \cdot 10^3$ to $20 \cdot 10^6$ K. HMI providing full-disk magnetic field observations of the Sun, these data help us to link different dynamics in the solar atmosphere from temperature to temperature and from location to location. The Interface Region Imaging Spectrograph (IRIS) [De Pontieu et al., 2014b] was launched in 2013. It is designed to study the solar chromosphere and TR, and it provides fruitful observations of the TR with unprecedented spatial and spectral resolutions. IRIS has the unique capability of acquiring both imaging and spectral data on the TR simultaneously, and it provides fruitful results on the dynamics of the TR (e.g. [De Pontieu et al., 2014a; Testa et al., 2014; Tian et al., 2014; Peter et al., 2014; Hansteen et al., 2014]. The above instruments have revealed the dynamic and complex nature of the TR, which presents many small-scale energetic events seen in imaging and spectroscopic data. In this paper, we review the observational facts of three types of such small-scale energetic events in the TR as transition region explosive events, UV bursts, and network jets.

TRANSITION REGION EXPLOSIVE EVENTS

In the 1970s, the Naval Research Laboratory launched three Black Brant sounding rockets with the High Resolution Telescope and Spectrograph (HRTS) experiment

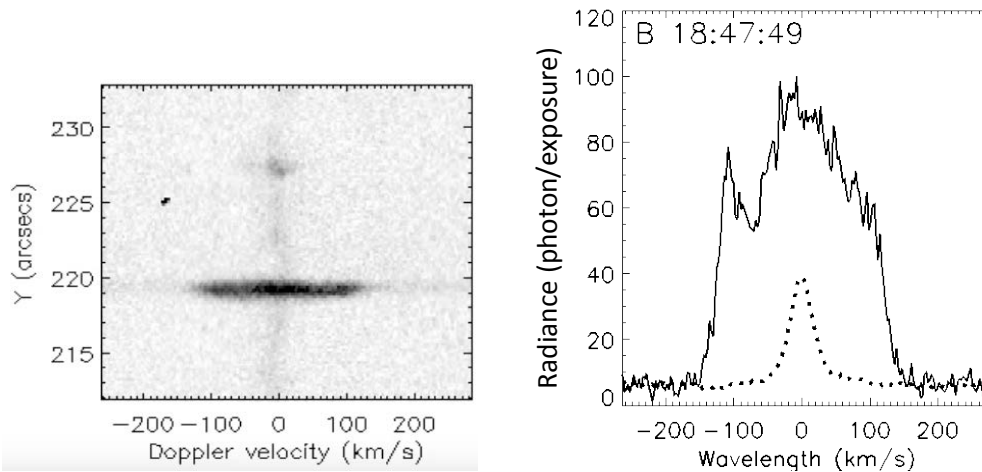


Figure 2. An example of Si IV spectral image of a transition region explosive event. The observations were taken by IRIS on June 27, 2013 as shown in Figure 5. Left panel — the event when it was crossing the slit. The image is shown in reversed-color scale, which presents strong emission in black and weak in white. The explosive event is presented at the position $Y=219\div 220$ arcsec. Right panel — the spectrum of the event (solid line). The normal spectrum is shown by dashed line (adapted from [Huang et al., 2014a])

onboard [Brueckner, Bartoe, 1983]. With a slit 900 arcsec long and 0.5 arcsec wide, the experiments provided spectral data over a wavelength range from 1175 to 1710 Å with a spectral resolution of 0.05 Å. The obtained spectra of Si IV ($8.0 \cdot 10^4$ K), C IV ($1.0 \cdot 10^5$ K), and O IV ($1.3 \cdot 10^5$ K) emitted from the TR exhibit profiles with strong wings extended to 50–250 km/s toward both short and long wavelengths. These events have 1–2 arcsec spatial extension along the slit, which has been taken as their spatial size. Their lifetime was however difficult to estimate because of the lack of imaging datasets. Huang et al. [2014a] have found that the brightening corresponding to a TREE could last for about one hour. Brueckner and Bartoe [1983] proposed that such spectra are emitted from turbulent events. While Dere et al. [1984] presented their study on the TR spectral data from HRTS, they called such events explosive type of events (i.e. explosive events). The term of transition region explosive events (TREES) started to be widely used since Dere et al. [1989] presented a statistical analysis of the HRTS data with detailed examples of such events. Figure 2 presents an example of TREE spectrum observed by IRIS. Although TREES are mostly identified in the TR emissions, some of them can also produce signatures in spectra from lower temperatures, such as C I, C II, O I, Lyman series, Mg II, etc (see, e.g., [Dere 1992; Madjarska, Doyle, 2002; Zhang et al., 2010; Huang et al., 2014b]).

According to the Doppler effect, the extended strong wings in the TR spectra indicate the existence of opposite-directional plasma flows in the source. Therefore, TREES have also been known as “bi-directional jets”. Such Doppler velocities can be generated by any mechanisms that can produce bi-directional flows within the space of the pixel size. For example, swirling jets suggest plasma motion along magnetic field lines [Curdt, Tian, 2011] and bi-directional flows along two close flux tubes [Alexander et al., 2013]. Because TREES normally also exhibit an enhancement of radiation, their observations also match the picture of magnetic recon-

nection, which predicts bi-directional outflows and energy release (see [Priest, Forbes, 2000; Priest, 2014] for theory and [Su et al., 2013; Yang et al., 2015; Sun et al., 2015; Li et al., 2016, 2018a; Xue et al., 2016; Huang et al., 2018] for real examples in the solar atmosphere). With this clue, magnetic reconnection is thought to be one of the mechanisms for producing TREES [Dere et al., 1989, 1991; Innes et al., 1997, 2015; Li, 2019], and has been evidenced by observations from instruments such as SUMER and IRIS (Innes et al., 1997; Huang et al., 2014b, 2018). Figure 3 shows spectral examples of TREES as seen by SUMER raster scan. It indicates that the spectra are one wing enhanced on one side, other wing enhanced on the other side, and both wing enhanced in the middle of the TREES. These observations provide a solid evidence for the magnetic reconnection scenario as shown in Figure 4.

The magnetic reconnection scenario has been indirectly supported by observations of TREE associated phenomena. TREES are usually related to mixed polarities of magnetic features at boundaries of supergranulation cells [Brueckner et al., 1988; Dere et al., 1989, 1991; Porter, Dere, 1991; Chae et al., 1998a; Madjarska, Doyle, 2003; Teriaca et al., 2004; Ning et al., 2004; Muglach, 2008]. Moreover, Chae et al. [1998a] have found that more than 60 % of TREES are associated with magnetic cancellation, a phenomenon usually occurring after magnetic reconnection [Zwaan 1987; van Ballegoijen, Martens, 1989; Huang et al., 2012, 2014a, b; Mou et al., 2016].

Moreover, many studies have focused on what dynamic features in the solar atmosphere can produce TREE spectra, and they have been found to be signatures of siphon-flows in small-scale loops [Teriaca et al., 2004], spicules [Wilhelm, 2000], chromospheric up-flows [Chae et al., 1998b], surges [Madjarska et al., 2009], swirling jets [Curdt, Tian, 2011], and coronal brightenings and jets [Madjarska et al., 2012; Huang et al., 2012; Li et al., 2018b]. These observations suggest that TREES can be triggered in ways other than magnetic reconnections.

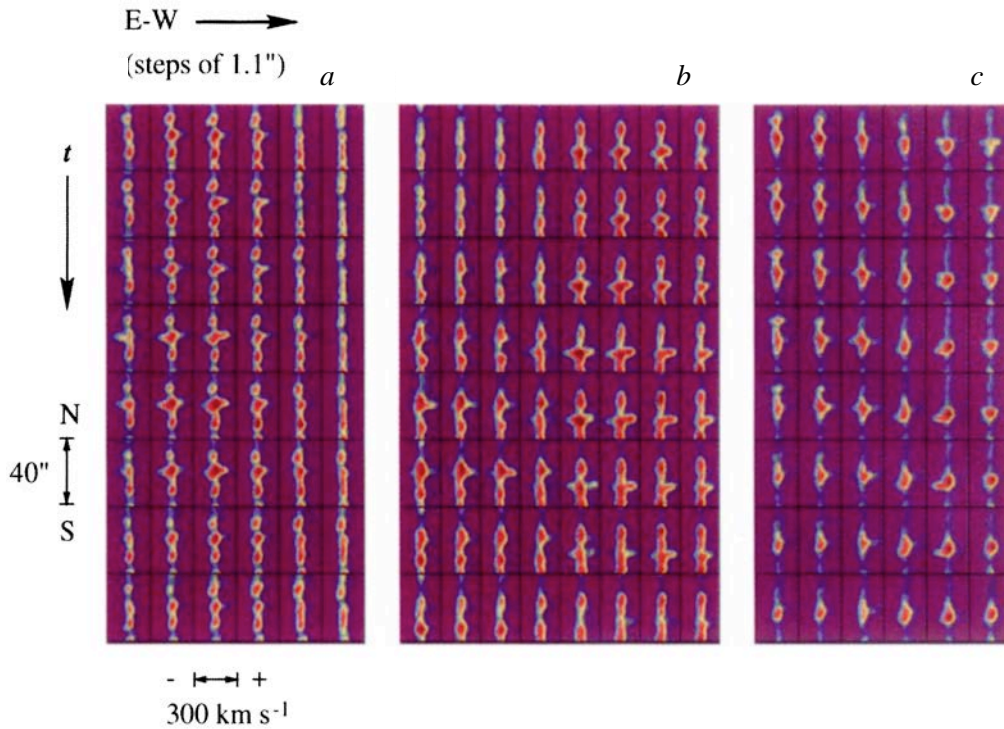


Figure 3. Spectra of TREEs as scanned by SUMER (adapted from [Innes et al., 1997])

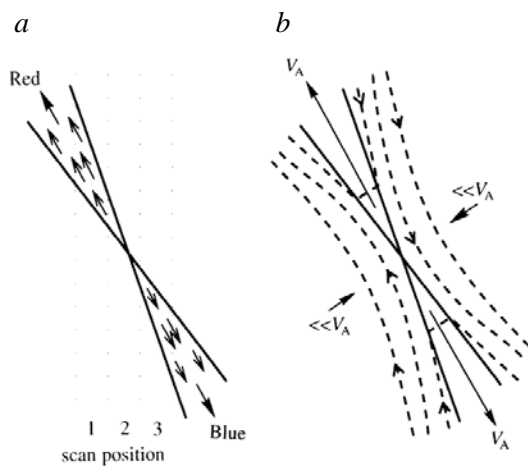


Figure 4. Magnetic reconnection scenario for TREEs based on the observations shown in Figure 3. The spectrograph scanning different locations (1–3) of the reconnection site, the spectra could present blue-wing enhancement, both wing enhancement and red-wing enhancement (adapted from [Innes et al., 1997])

Recently, the very high-resolution imaging and spectral data that are simultaneously taken by IRIS have revealed more details of TREEs. Huang et al. [2014a] studied a TREE that was associated with a dot-like brightening at a coronal-hole boundary, and they found that the TREE was associated with interactions between small-scale loops and small-scale jets (see Figure 5). The authors suggested that the interaction between loops could lead to magnetic reconnection that is responsible for the TREE. The connection between the loop-loop interaction and TREEs has also been confirmed by Huang et al. [2015]. Huang et al. [2017] have studied 24 TREEs found in active regions, aiming at finding out which phenomena in the chromosphere and

transition region are associated with TREEs. They have established that 16 TREEs are associated with stationary loop-brightenings without any response in chromospheric H α spectra; six are associated with propagating loop brightenings, in which two events are associated with dark jets observed in chromospheric H α lines; two are associated with brightenings in the conjunction region of multiple loops seen in the transition region. The study by Huang et al. [2017] clearly indicates the close relation between TREEs and dynamics of TR loops. Since the dynamics of loops could directly link to how mass and energy are transported in the solar atmosphere [Klimchuk, 2006], the interaction between two loop systems might result in magnetic reconnection that could exchange plasma between the two systems and generate heats therein [Huang et al., 2018]. Such loop interactions might generate TREE spectra (see [Huang et al., 2015]), and thus TREE could be a signal for searching magnetic reconnection between loops in the solar atmosphere. In the other conditions, a loop system might contain multiple threads that might be braided due to the random motions in the photosphere [Parker, 1983a]. The braiding magnetic structures could build up magnetic energy via small-scale magnetic reconnection events in the solar atmosphere [Parker, 1983b, 1988; Klimchuk, 2006; Schrijver, 2007; Cirtain et al., 2013]. Huang et al. [2018] have found that magnetic braidings of a transition region loop system can create plasma jets at the Alfvén speed with TREE spectra observed in their footpoints. The observations made by Huang et al. [2018] confirm the suggestion that magnetic braiding is also a possible mechanism for producing TREEs [Huang et al., 2017].

Based on these new observations, some modeling attempts have been put forward. Innes et al. [2015] have performed a two-dimensional simulation. They have found

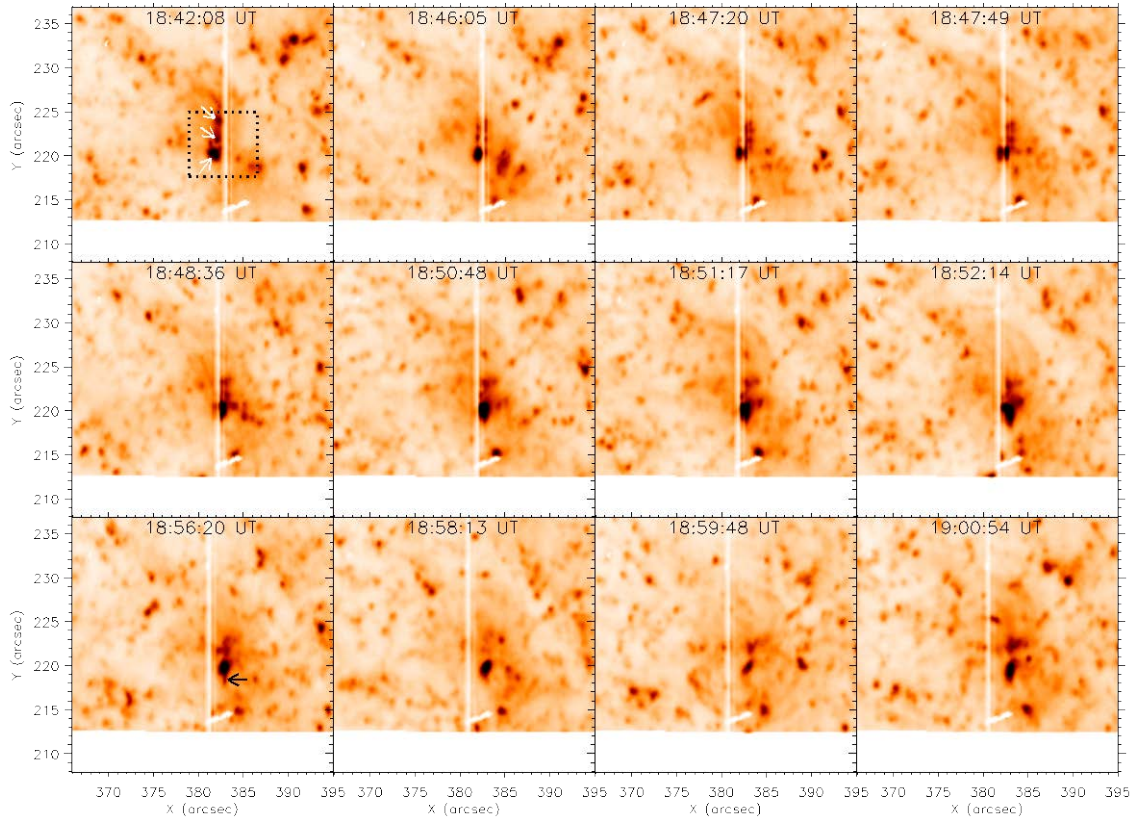


Figure 5. Temporal evolution of a TREE (pointed by arrows in the top-left panel) as seen in IRIS SJ 1330 Å. The images are shown in reversed color scale (adapted from [Huang et al., 2014a])

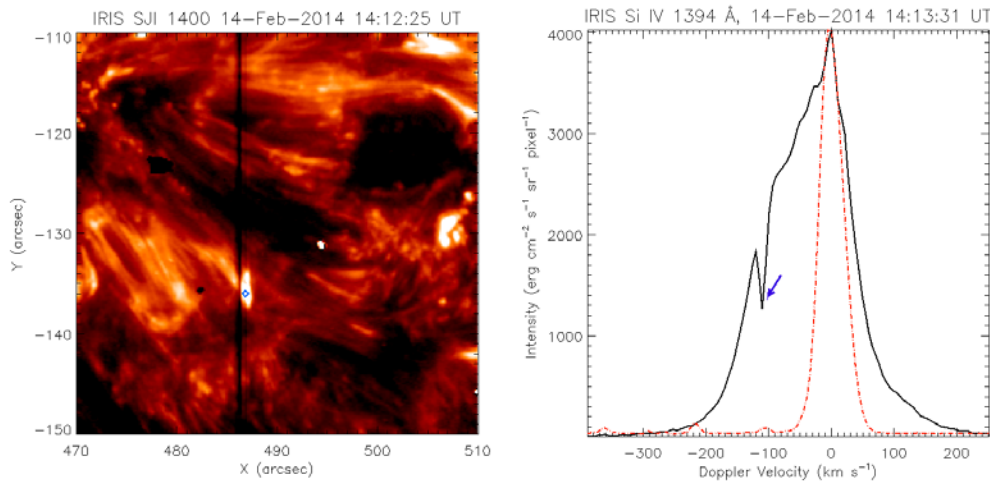


Figure 6. An example of IRIS bomb. The IRIS bomb is the compact bright feature marked by the blue diamond in the IRIS SJ 1400 Å image (left panel). The IRIS Si IV 1394 Å spectrum sampled from the location of the blue diamond is shown on the right panel (black solid line). The blue arrow points at the location where the absorption lines of Ni II are superimposed in the wing of the Si IV line. The red dashed-dotted line on the right panel is the referent spectrum that is taken from a quiet region and multiplied by 9 times. Courtesy of Mr. Zhenyong Hou

that TREE profiles can be reproduced with the multiple magnetic islands and acceleration sites that characterize the plasmoid instability if TREES are observational signatures of magnetic reconnection sites.

UV BURSTS IN THE TR

Loop interactions in the stages of flux emergence can emit spectra with characteristics more than those of the TREES. Peter et al. [2014] have reported on a new

type of small-scale events seen in IRIS transition region observations, named “hot explosions” (see, e.g., Figure 6).

In the IRIS imaging data, these events are transient compact bright features with intensity 100–1000 times higher than the median background and a size of a few arcseconds. These events are observed in TR Si IV lines showing very broadened profiles together with clear signatures of chromospheric absorption lines (e.g., Fe II and Ni II blended in the blue wing of Si IV 1394 Å; and Fe II, in the red wing of Si IV 1403 Å). These events are

also called “IRIS bombs” [Vissers et al., 2015; Tian et al., 2016]. Since both IRIS bombs and TREEs have similar size and emit similar broad TR spectra, the blends of cold absorption lines are the main differences that distinguish these events from TREEs. An ISSI group had been assembled aiming at studying these events, and they introduced a term “UV bursts” in the TR to describe small, intense, transient brightenings in ultraviolet images of solar active regions including IRIS bombs as described above (see the definition in [Young et al., 2018]).

Many UV bursts are believed to be the result of magnetic reconnection at the early stage of flux emergence, when the flux tube might show a serpentine (or U-loop) geometry in the lower solar atmosphere [Fan, 2001; Pariat et al., 2004, 2009; Cheung et al., 2007, 2008; Archontis et al., 2009; Cheung, Isobe, 2014; Huang et al., 2018]. UV bursts in the TR occur frequently at the early stage of flux emergence (see, e.g. [Peter et al., 2014; Toriumi et al., 2017; Zhao et al., 2017; Tian et al., 2018a] and rarely in the late phase of flux emergence (see, e.g. [Guglielmino et al., 2018; Huang, 2018]). The UV bursts reported in Peter et al. [2014] were first suggested to be caused by magnetic reconnections in the photosphere, which heat plasma to the TR temperature. Judge [2015] performed an independent analysis and suggested that these events were formed in the low to middle chromosphere or above. While Ellerman bombs are phenomena that are well known to be signatures of magnetic reconnections in the photosphere [Ellerman, 1917; Vissers et al., 2015; Reid et al., 2016; Fang et al., 2006; Hong et al., 2014; Nelson et al., 2013, 2015, 2016; Fang et al., 2017; Ni et al., 2016, 2018a], Tian et al. [2016] have crosschecked IRIS observations and H α images of 10 TR UV burst events from the Chinese New Vacuum Solar Telescope [Xu et al., 2014; Liu et al., 2014] and have found that 6 of the 10 events are associated with Ellerman bombs in the photosphere and the rest 4 are not. This indicates that some UV bursts in the TR are signatures of magnetic reconnections in the photosphere that heat plasma up to at least the TR temperature [Kim et al., 2015; Tian et al., 2016]. Such a conclusion has also been confirmed by [Nelson et al., 2017], who established that some Ellerman bombs in the quiet Sun also have responses in the IRIS Si IV lines. It is worth mentioning that the blends of the chromospheric absorption lines in the TR emission lines might also be the results of cold plasma (such as filaments, mini-filaments) at the foreground [Vissers et al., 2015; Huang et al., 2017], and thus it is not complete criteria for this species of UV bursts.

TR UV bursts are also diverse. Hou et al. [2016] have reported on a different type of UV bursts, so-called “Narrow-line-width UV bursts” (or “NUBs” for short). NUBs are also compact brightenings but showing very different characteristics in the TR spectra. The TR spectra (Si IV) of NUBs are regularly Gaussian but extremely narrow (less than half of that from a normal region). NUBs are normally found locating above sunspots and might consist of multiple bright cores.

With these observations, some simulation experiments have tried to explain whether magnetic reconnection in the photosphere could heat plasma to transition

region temperatures. Ni et al. [2018a, b] have found that initially weakly ionized plasmas can become fully ionized within the reconnection region and the current sheet can be strongly heated to above $2.5 \cdot 10^4$ K in the cases that the reconnecting magnetic field is in excess of a kilogauss and the plasma β is lower than 0.0145. For more details of UV bursts, we refer to a review paper by Young et al. [2018].

JETS IN THE TR

Besides the energetic events outstanding from the brightness and spectra, jets also frequently occur in the TR. The IRIS observations have revealed that jets are ubiquitous in the TR network regions and so-called network jets [Tian et al., 2014], and see Figure 7 for examples). Network jets in coronal hole regions are small-scale with widths of ≤ 300 km, lifetimes 20–80 s, and speeds 80–250 km/s (see Figure 8 and [Tian et al., 2014; Qi et al., 2019]). Network jets have also been found in the quiet-sun region. These jets are typically slower and shorter than those in coronal holes [Narang et al., 2016; Kayshap et al., 2018]. Some of the network jets should be powered by magnetic reconnection and many could be heated to at least 10^5 K [Tian et al., 2014]. The appearances, sizes, lifetimes, and speeds of the network jets have a large overlap with that of spicules (see characteristics of spicules in, e.g., [Xia et al., 2005; De Pontieu et al., 2007; Zhang et al., 2012; Pereira et al., 2012, 2014]). Since both network jets and spicules originate in the network regions, it suggests that at least some of spicules in the chromosphere should be heated to the TR temperature and shown as network jets. A recent simulation by Yang et al. [2018] based on magnetic reconnection driven by a combination of magnetic flux emergence and horizontal advection shows that both fast warm jets (much similar to the network jets) and slow cool jets (much similar to classical spicules) could be generated at a time. Recent analytical studies and simulations on magnetic reconnections driven by magnetic cancellation [Priest et al., 2018; Syntelis et al., 2019] could release

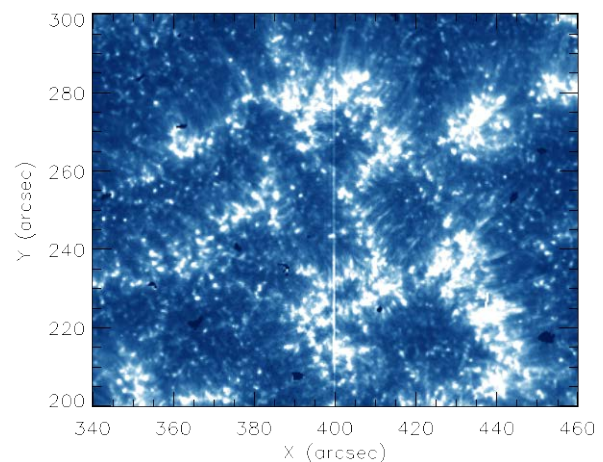


Figure 7. A network region observed in the IRIS SJ 1330 Å passband on December 4, 2015 at 01:29 UT. The compact brightenings are organized to be the so-called “network lanes”. The elongated thin features extending from the bright network lanes are network jets. Courtesy of Ms. Youqian Qi

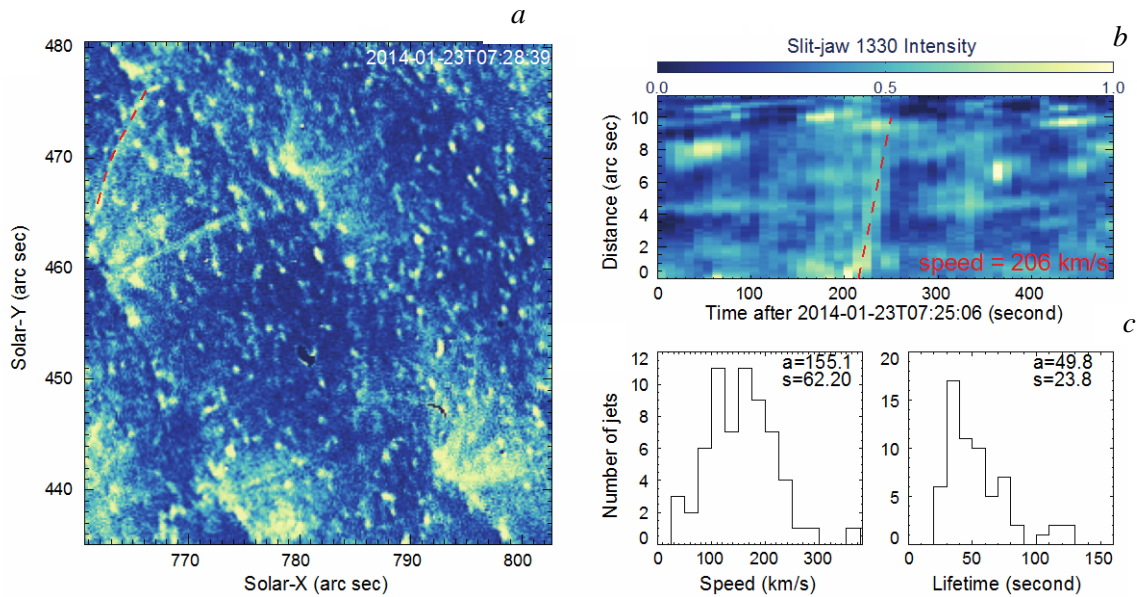


Figure 8. Network and jets seen in IRIS SJ 1330 Å (a), time-distance diagram of a jet (b) and statistic histograms of speeds and lifetimes of network jets (c) (adapted from [Tian et al., 2014])

energy at different height thus produce energetic events at different temperatures.

Chen et al. [2019] have carried out a statistic analysis aiming at searching for the connection between TREEs and network jets. They have found that (1) TREEs with double peaks or enhancements in both wings appear to be located at either the footpoints of network jets or transient compact brightenings. These TREEs are most likely originated from magnetic reconnection; (2) TREEs with enhancements only at the blue wing are mainly located above network jets along their propagation directions. These TREEs clearly result from the superposition of high-speed network jets onto the TR background; (3) TREEs showing enhancement only at the red wing of the line are generally located around the footpoints of network jets, likely caused by the superposition of reconnection downflows onto the background emission. Their analysis suggests that some TREEs are related to the birth or propagation of network jets, and that other TREEs are not connected to network jets.

Furthermore, Qi et al. [2019] have statistically studied network jets and have found that network jets in the footpoints of strong coronal plumes are more dynamic (higher and faster) than those from regions without clear coronal plumes. This suggests that more dynamic network jets might somehow connect to hotter features in the corona. It has been reported that spicules [De Pontieu et al., 2011; Pant et al., 2015; Jiao et al., 2015, 2016; Samanta et al., 2015] and/or shocks [Hou et al., 2018] could be the possible sources of the propagating disturbances in the coronal plumes. It is likely that some of the network jets should also be heated to coronal temperatures.

Besides network jets, there are also some other types of small-scale jets in various regions of the TR. Huang et al. [2018] have reported on very fine jets occurring during the eruption of a braiding structure that contains a quasi-open field. The jets have speeds higher than

300 km/s, which are comparable to the Alfvén speed in the eruptive region. The observations shown in Huang et al. [2018] also demonstrate that these jets result from magnetic reconnections due to magnetic braiding. Tian et al. [2018b] have reported on high-speed jets found above the light bridge of sunspots. These jets could be clearly seen in images taken at offbands of H α , and they also show signatures of bi-directional flows with speeds as large as 200 km/s in TR spectral observations. These jets also exhibit clear inverted Y shapes, which clearly indicate magnetic reconnection between close and open magnetic fields.

CONCLUSION

The solar transition region (TR) is the crucial interface region between the chromosphere and corona. It is normally defined in the temperature regime from about $2 \cdot 10^4$ K to about 10^6 K. The space-born instruments have opened windows for studying this region and many dynamic phenomena have been discovered. This paper reviews a few types of small-scale energetic events occurring in the TR. These events are defined in either spectral or imaging data. Multi-instrument observations available recently have given opportunities to view the same phenomenon in both spectral and imaging data, and it has given more insights to the phenomenon. The observations of these small-scale energetic events indicate the extremely dynamic and complex nature of the TR and they demonstrate that the dynamics of the TR is crucial in the processes of mass and energy transportation between the solar chromosphere and the corona.

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